CURRENT GROUND-BASED DETECTORS

And their likely evolution to 2040

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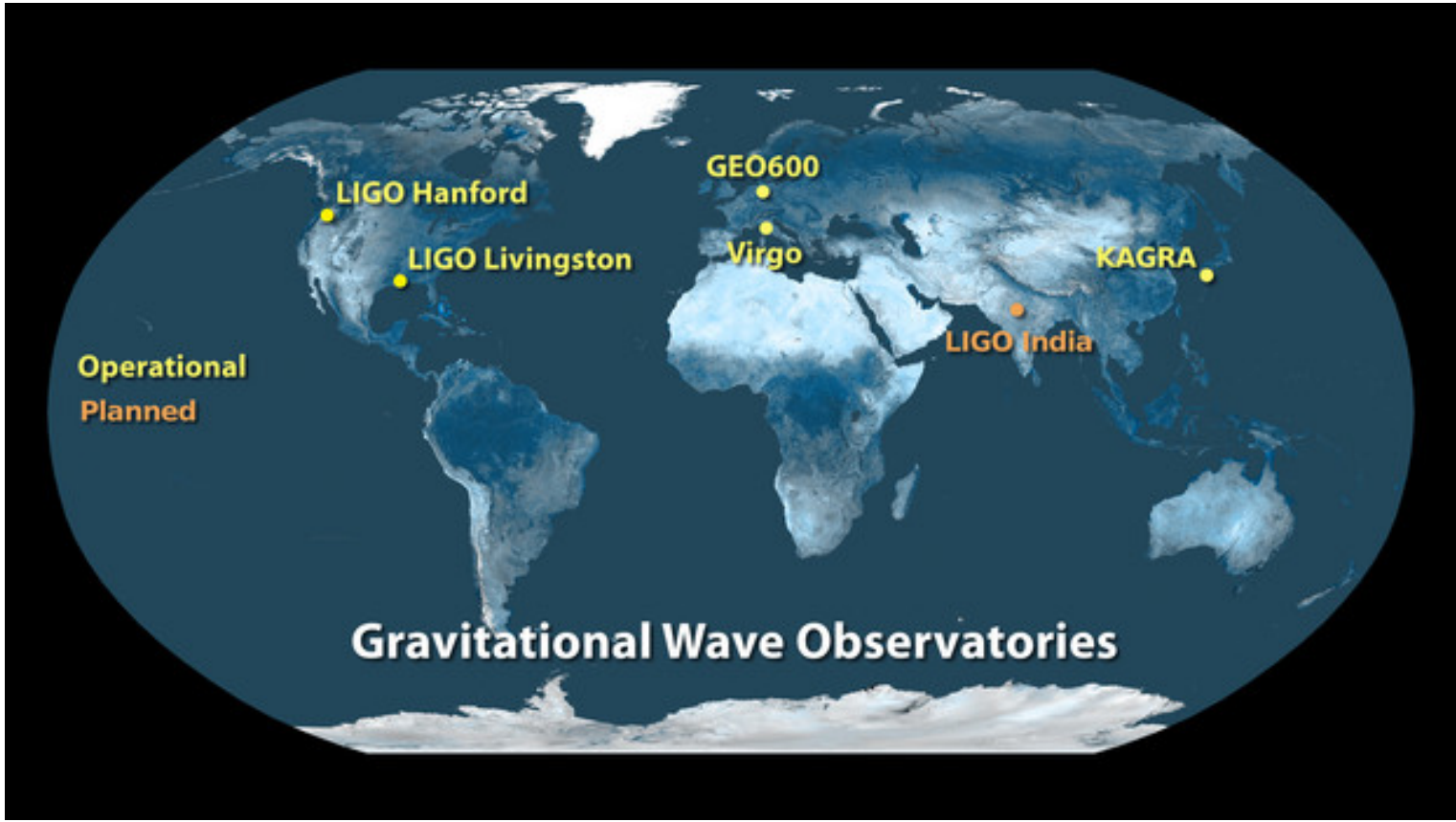
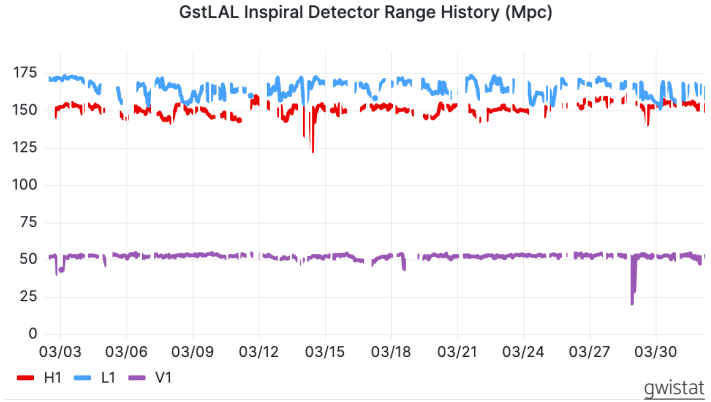
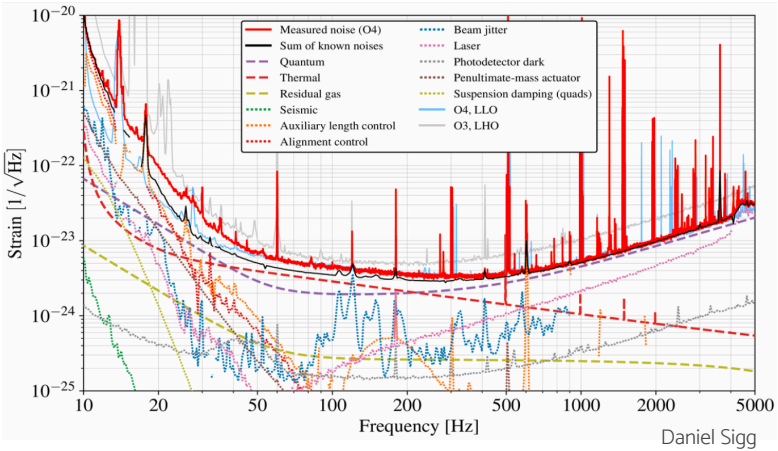
10 Years to LISA, JPL, 1st April 2025

WHERE ARE WE NOW?

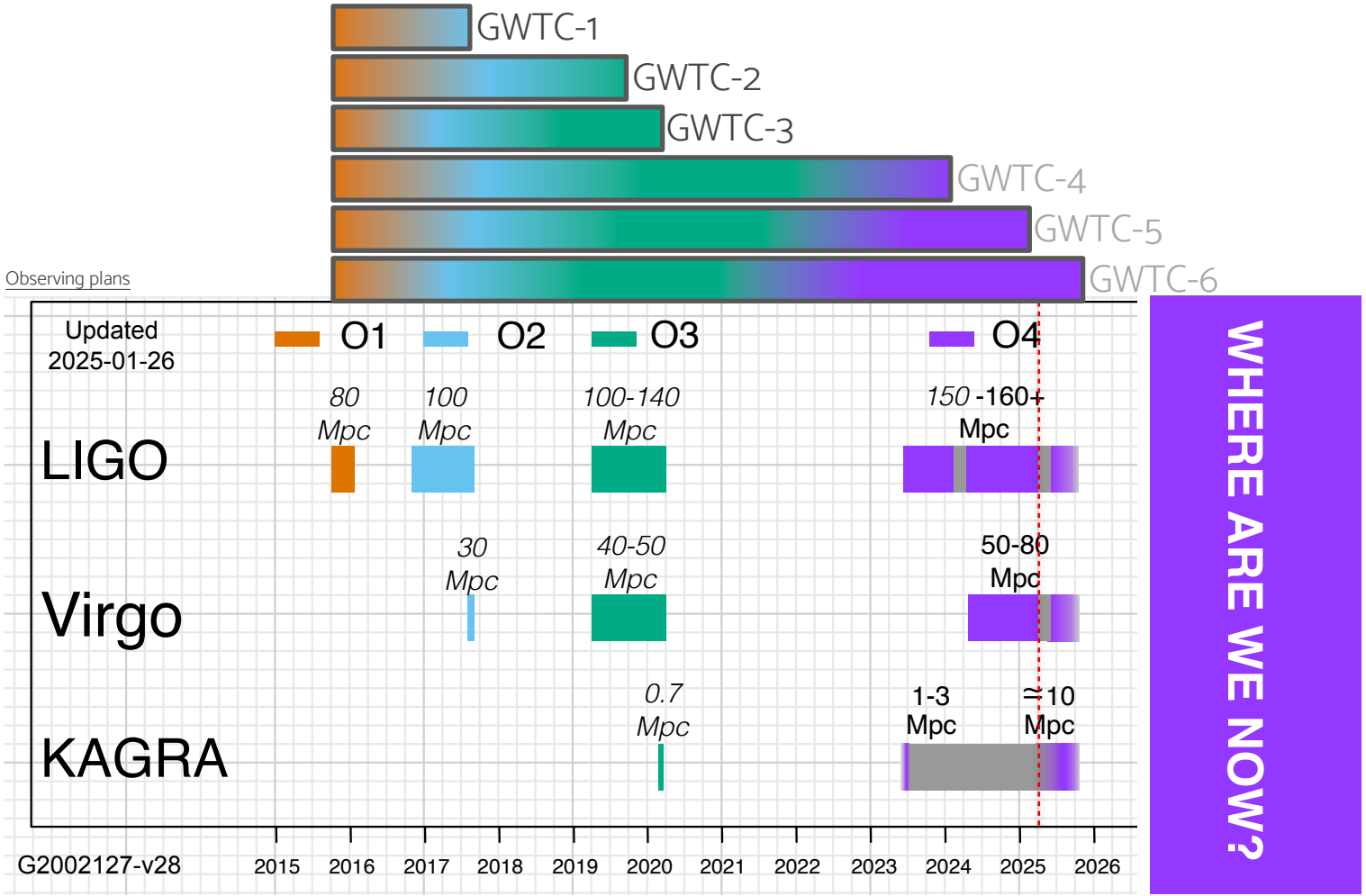


WHERE ARE WE NOW?

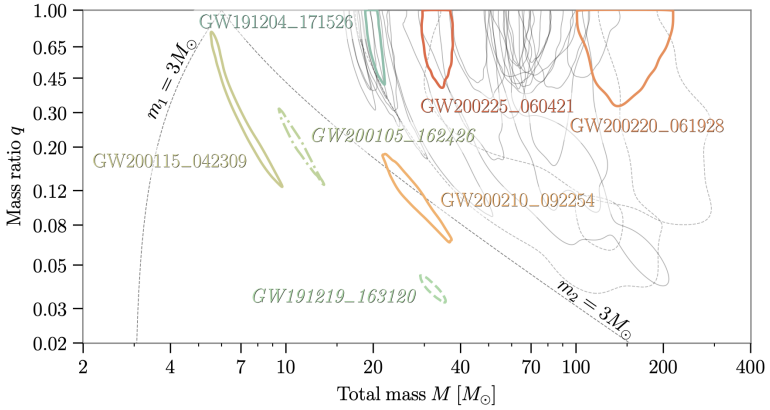


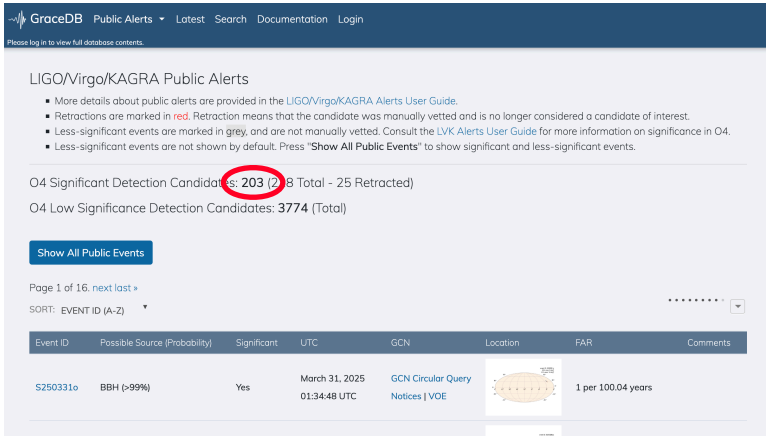


- ALIGO detectors: dual-recycled Michelson interferometers, 4km arms, 40kg test masses, with frequency-dependent squeezing
- AdVirgo: 2km arms
- Most downtime aligned to maximise 3-detector uptime
- KAGRA: underground for lower seismic noise, cryogenic
- GEO600: used for developing and testing, and Astrowatch

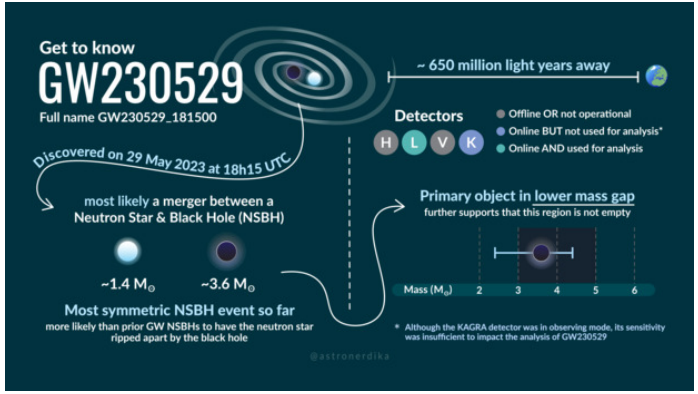

GWTC-1
GWTC-2
GWTC-3
GWTC-4
GWTC-5
GWTC-6
Observing plans
Updated
2025-01-26
O1
O2
O3
O4
LIGO
80 Mpc
100 Mpc
100-140 Mpc
150 -160+ Mpc
Virgo
30 Mpc
40-50 Mpc
50-80 Mpc
KAGRA
0.7 Mpc
1-3 Mpc
≃10 Mpc
G2002127-v28
2015
2016
2017
2018
2019
2020
2021
2022
2023
2024
2025
2026
WHERE ARE WE NOW?

WHERE ARE WE NOW?

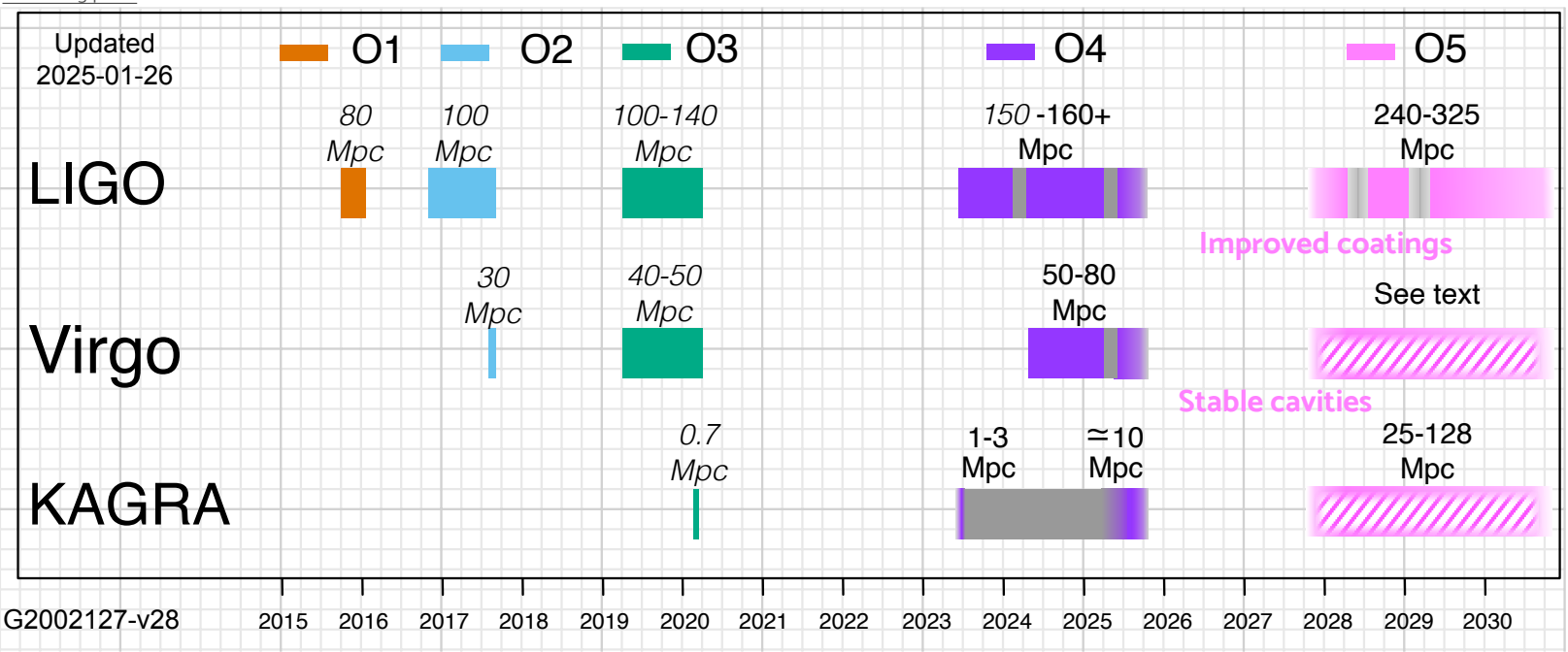




- GWTC-3: 90 confident detections
- Majority BBH mergers
- ~ a few BNS mergers, notably GW170817
- ~ five NSBH mergers, notably GW200105 and GW200115
- Some systems with masses between 3 and 5 $M_\odot$, GW190814 and GW230529



Observing plans

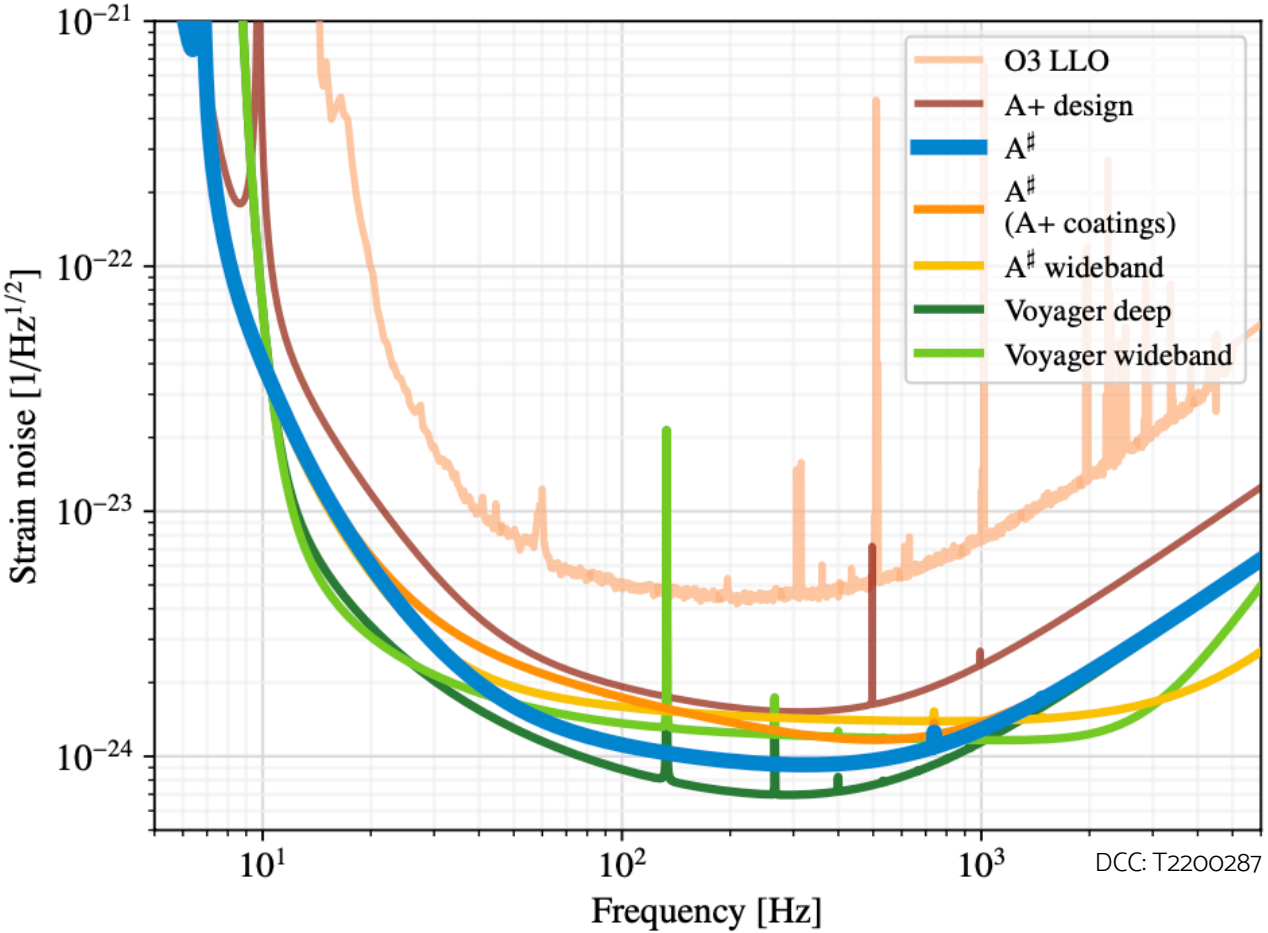


TOWARDS A+

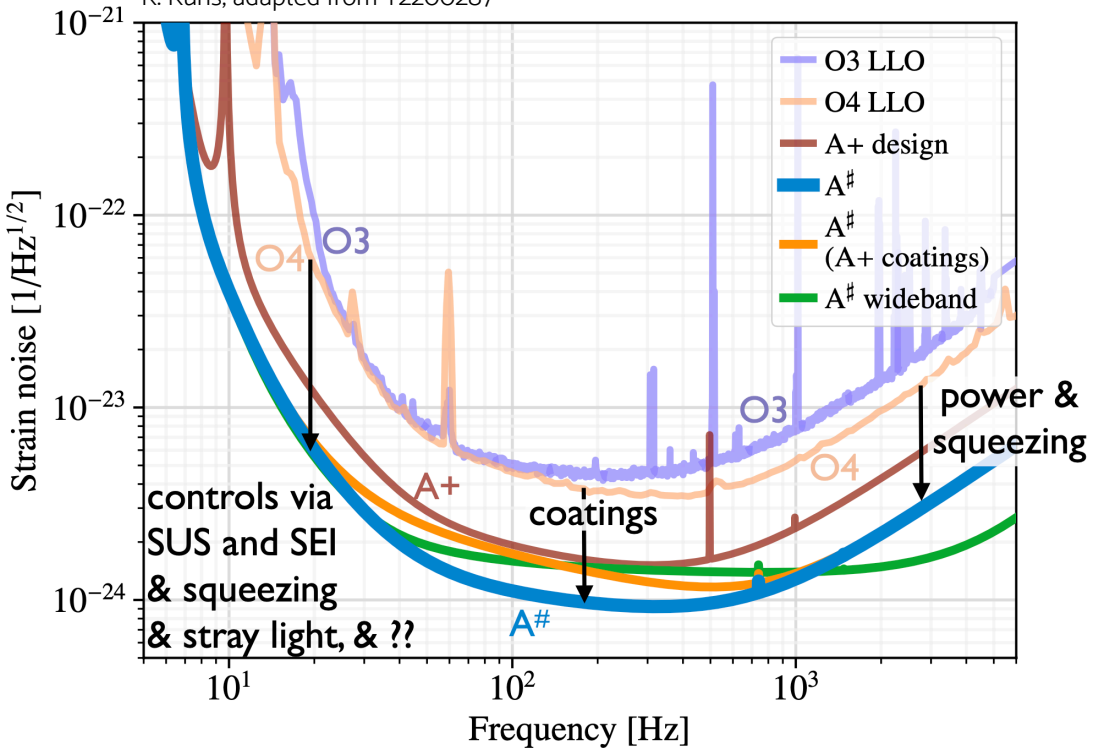
INTO THE NEXT DECADE: A#

A bridge between A+ and next-generation ground-based instruments, and a testing ground for new technologies.

See also: Virgo_nEXT



K. Kuns, adapted from T2200287



LOW FREQUENCIES

- 100kg optics on upgraded suspensions
- Use suspension fibres at higher stress

MID FREQUENCIES

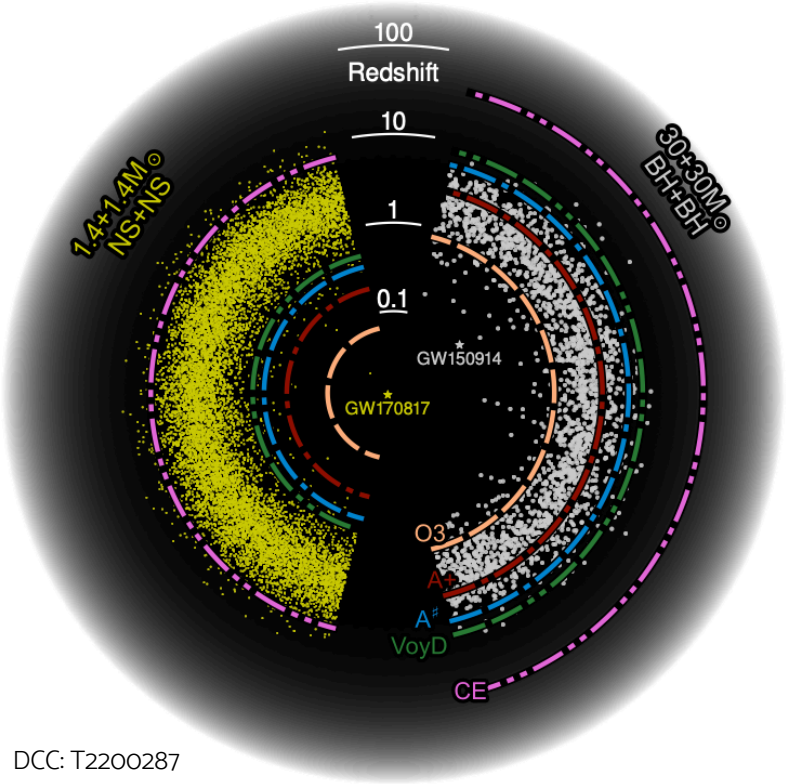
- Improve coatings for lower thermal noise

HIGH FREQUENCIES

- Increase circulating power to 1.5MW
- Increase squeezing up to 10dB

A# KEY UPGRADES

A# DETECTION ESTIMATES



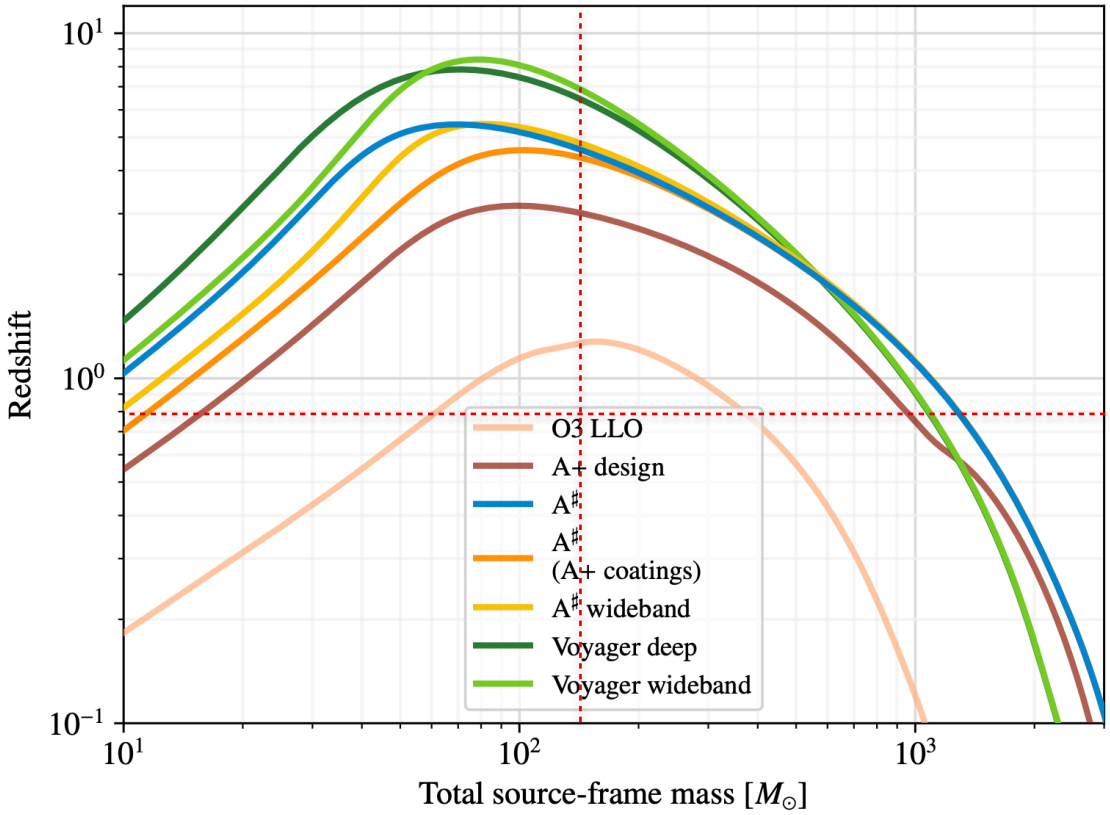
DCC: T2200287

- BNS and NSBH rate in A# more than 4 times that of A+
- Even with A+ coatings, A# still doubles BNS and NSBH rates
- BBH rate is almost 3 times with A# compared to A+

Configuration	Annual Detections		
	BNS	NSBH	BBH
A+	135^{+172}_{-78}	24^{+34}_{-16}	740^{+940}_{-420}
A♯	630^{+790}_{-350}	100^{+128}_{-58}	2100^{+2600}_{-1100}
A♯ (A+ coatings)	260^{+320}_{-140}	45^{+60}_{-27}	1150^{+1450}_{-640}
A♯ Wideband (A+ coatings)	200^{+250}_{-110}	40^{+54}_{-25}	970^{+1220}_{-540}

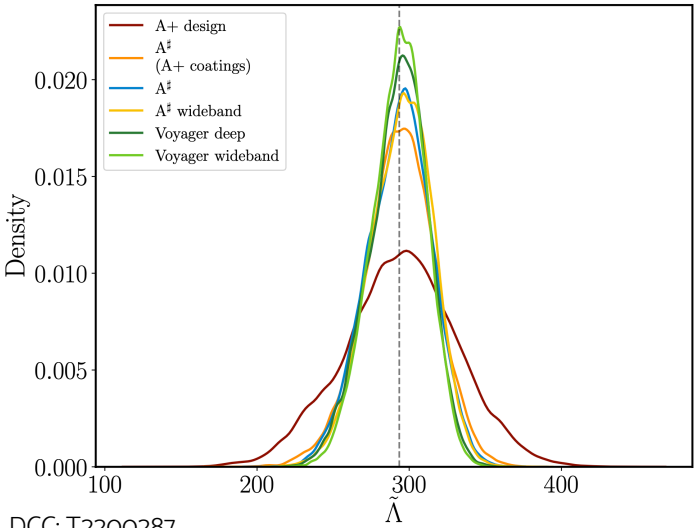
MASS AND REDSHIFT HORIZONS

- A+ will be able to see GW190521-like binaries out to z=3, and A# out to z=5
- Could see ten times heavier GW190521 at the same redshift



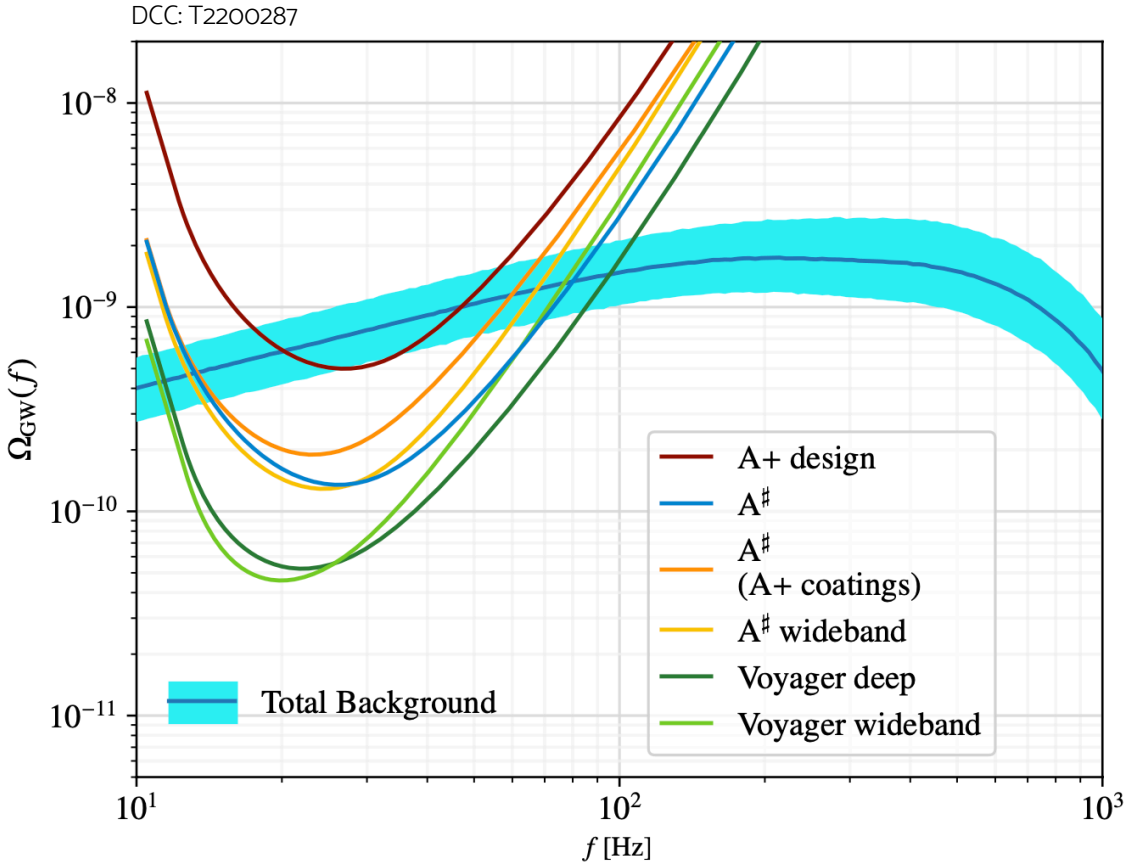
NS SCIENCE

Configuration	Range [Mpc] BNS	BBH	t_{early}[min]	z_{max}	Post-Merger $\rho_{\mathrm{pm}}^{(10)}$	$\rho_{\mathrm{pm}}^{(\mathrm{max})}$
O3 LLO	130	1200	0.3	1.3	0.4	0.6
July 2022 LLO	120	1200	0.5	1.5	0.3	0.5
A+	350	2600	2.7	3.2	1.4	2.0
A+ Wideband	290	2300	3.7	3.5	2.2	2.6
A♯	600	3700	6.2	5.4	2.7	3.7
A♯ (A+ coatings)	440	3000	6.1	4.6	2.7	3.4
A♯ Wideband	490	3300	6.8	5.5	4.8	5.6



DCC: T2200287

- BNS range around 4 times further than current
- Low frequency sensitivity improvements lead to BNS early warning of six minutes
- High frequency improvements make BNS post-merger detections a tantalising possibility
- Can infer tidal parameter significantly better



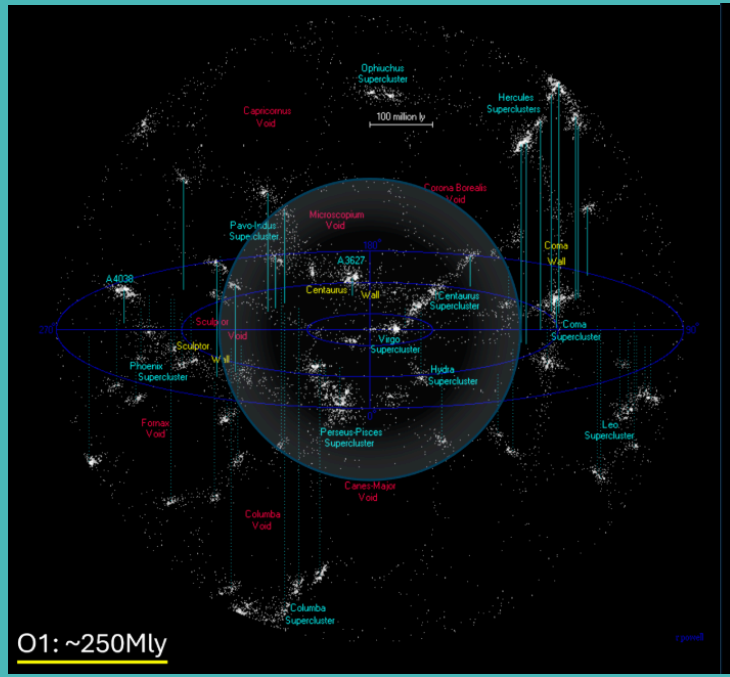
- If not already detected by O5, any post-O5 upgrade will allow for the detection of an unresolved background of binary mergers

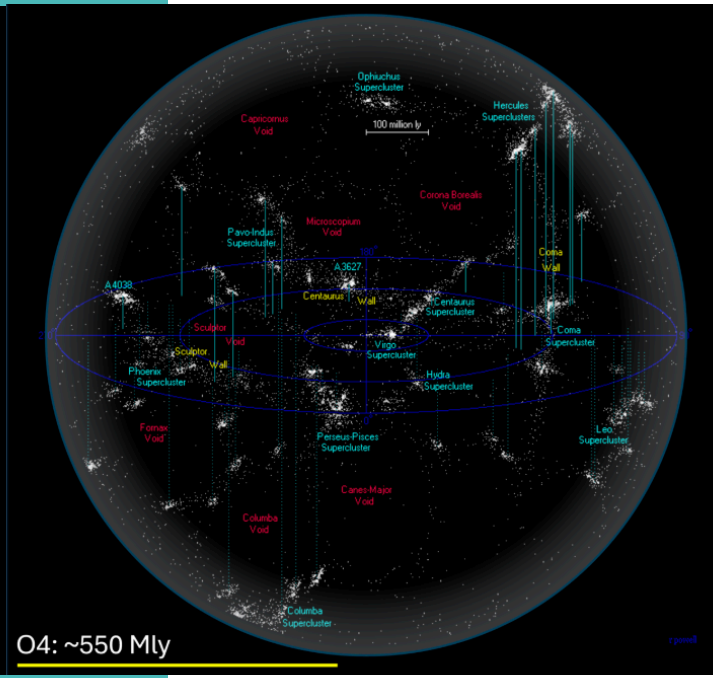
UNRESOLVED BACKGROUND

MIT/Caltech/LIGO Lab/Kim Burtnyk

SUMMARY

- There have been huge sensitivity improvements since O1, and correspondingly many more detections
- A+ is only three years away, and A# in the early 2030s will provide significant science benefits through improved coatings, suspensions, larger test masses and higher laser power
- Upgrades will lead to many more detections, including heavier BBHs, better NS constraints and an unresolved background of merging binaries





THANK YOU!